

Variable Star Bulletin

Outburst of RX J0575.0+6306

Tsutomu Watanabe

117 Shirao dormitory , 1414 Onakazato , Shizuoka , 418-0044 Japan

E-mail : jcc00212@nifty.ne.jp

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1 Observations

The ROSAT source RX J0757.0+6306 was optically identified with a cataclysmic variable located at R.A. 07h 57m 01.3s Dec. $+63^{\circ} 06' 01''$ (J2000.0) , having a 75 min orbital period , and showing a spectrum resembling that of WZ Sge (Tovmassian et al. 1997, 1998a). VSOLJ members carried out 86 visual observations between 1997 Oct. and 1998 Jun.: T.Watanabe (58 obs.) , H.Itoh (26 obs.) , S.Takahashi (1 obs.) and M.Hiraga (1 obs.). The Light curve is shown in Figure 1. One short outburst was detected on 1997 Dec 12 at a magnitude of 12.9 by H.Itoh. Subsequent confirmation by other observers was prevented by the waxing moonlight and the low transparency characteristic to the season. He later confirmed observed position was truly correct. The duration the object was brighter than 13.0 magnitude is shorter than 3 days. Observation data near the outburst are listed in Table 1. Considering the magnitude and the duration , the outburst seems to be a normal outburst.

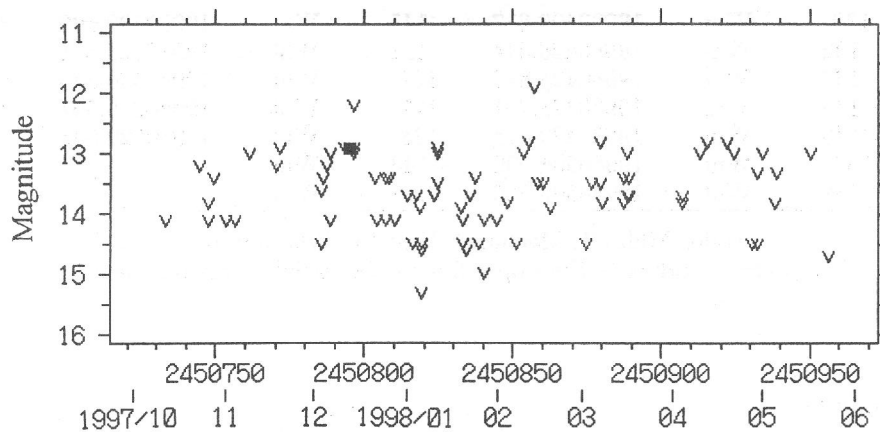


Figure 1 : Light curve of RX J0757.0+6306 between 1997 Oct. and 1998 Jun.

Table 1. Observation data near the outburst

Date(UT)	Mag. ^a	Observer
19971205.860	<131	Ioh
19971210.663	<130	Wnt
19971212.866	129	Ioh
19971213.549	<123	Ioh
19971213.549	<123	Ioh
19971213.578	<130	Ths
19971213.669	<131	Ioh
19971220.797	<135	Wnt

^a The pre- $x <$ indicates the upper limit (the variable was not seen)
code: Ioh (H. Ito), Ths (S. Takahashi), Wnt (T. Watanabe)

2 Acknowledgments

I wish to thank Dr.Taichi Kato (Kyoto University) for his helpful suggestions and reading the manuscript.

3 Remarks

This outburst detection was announced in [vsnet-obs 8928], which was referred to and discussed in Tovmassian et al. (1998b).

References

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Light Curve of SBS1017+533

Tsutomu Watanabe

117 Shirao dormitory , 1414 Oonakazato , Shizuoka , 418-0044 Japan

E-mail : jcc00212@nifty.ne.jp

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1 Observations

SBS1017+533 (R.A.10h 20m 30s Dec.+53° 02' 07" (2000)) is a cataclysmic variable discovered by the SBS survey (Balayan 1997). The star was identified with the ROSAT X-ray source RX J1020.4+5304 (Bade N. 1998 ,independently Kato 1998 and Garnavich 1998) The first outburst was discovered by Garnavich (1998) on 1998 Jan. 27.5. Superhumps with a period of 0.069 d were detected (Vanmunster 1998a, Nogami 1998), confirming the UGSU

Table 1. Magnitude estimates of SBS1017+533.

Date(UT)	Mag. ^a	Observer	Date(UT)	Mag. ^a	Observer	Date(UT)	Mag. ^a	Observer
19980203.643	130	Wnt	19980322.689	<132	Mhh	19980520.528	<125	Mhh
19980203.643	130	Wnt	19980403.617	<145	Wnt	19980522.600	<143	Wnt
19980215.646	<132	Wnt	19980426.544	<151	Wnt	19980523.501	<145	Wnt
19980217.656	<137	Wnt	19980428.617	127	Wnt	19980530.549	<143	Wnt
19980301.691	<145	Wnt	19980428.701	127	Wnt	19980615.597	<135	Wnt
19980303.653	<143	Wnt	19980429.515	138	Wnt	19980620.533	<143	Wnt
19980315.556	135	Wnt	19980504.596	<143	Wnt			
19980316.596	<138	Wnt	19980505.526	<137	Wnt			

code: Mhh (H. Maehara), Wnt (T. Watanabe)

^a The pre-x < indicates the upper limit (the variable was not seen)

type. After this outburst, VSOLJ members detected additional 2 normal outbursts (total 3 outbursts) by visual observations between 1998 Feb. and 1998 Jun. The light curve is shown in Figure 1. and the original data are given in Table 1. Both normal outbursts were confirmed by other observers (Kinunnen 1998 , Poyner 1998 and Vanmunster 1998b) As shown in Figure 1, the mean cycle length of outbursts is about 45 days. Further observations are strongly encouraged.

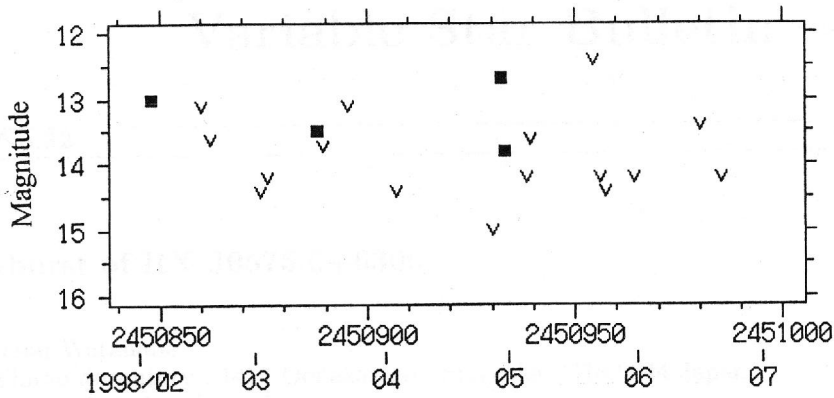


Figure 1 : Light curve of SBS1017+533.

2 Acknowledgments

I wish to thank Dr. Taichi Kato (Kyoto University) for his helpful suggestions and reading the manuscript.

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Variable Star Observers League in Japan
National Science Museum
Ueno Park Taito-ku Tokyo 110 Japan

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